

Black Hole Evaporation by Thermal Bath Removal *

J.Cruz[†] and J.Navarro-Salas[‡]

Departamento de Física Teórica and IFIC, Centro Mixto Universidad de Valencia-CSIC.
Facultad de Física, Universidad de Valencia, Burjassot-46100, Valencia, Spain.

Abstract

We study the evaporation process of 2D black holes in thermal equilibrium when the incoming radiation is turned off. Our analysis is based on two different classes of 2D dilaton gravity models which are exactly solvable in the semiclassical approximation including back-reaction. We consider a one parameter family of models interpolating between the Russo-Susskind-Thorlacius and Bose-Parker-Peleg models. We find that the end-state geometry is the same as the one coming from an evaporating black hole formed by gravitational collapse. We also study the quantum evolution of black holes arising in a model with classical action $S = \frac{1}{2\pi} \int d^2x \sqrt{-g} (R\phi + 4\lambda^2 e^{\beta\phi})$. The black hole temperature is proportional to the mass and the exact semiclassical solution indicates that these black holes never disappear completely.

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[†]CRUZ@LIE.UV.ES

[‡]JNAVARRO@LIE.UV.ES

1 Introduction

One of the main features of 1+1 dilaton gravity [1] (see the reviews [2] [3] [4]) is that the formation and semiclassical evaporation of a black hole can be studied in an exact analytical setting. The Russo-Susskind-Thorlacius (RST) model [5] predicts, with a natural choice of boundary conditions, an evaporation process ending in the linear dilaton vacuum in a finite proper time. Recently, Bose-Parker-Peleg [6] have proposed a different one-loop quantum modification of the CGHS model [1] describing a semiclassical evaporation with a semi-infinite throat as the remnant geometry. In both models the matching between the evaporating black hole solution and the static end-state geometry produces a thunderpop carrying a small amount of negative energy. The same physical picture of black hole evaporation exhibits the one-parameter family of models interpolating between the RST and BPP models [7]. In this paper we consider a different scenario to study the process of black hole evaporation in these theories. Instead of analyzing the formation and subsequent evaporation of a black hole we shall consider the dynamical evolution of a black hole in thermal equilibrium when the incoming flux decreases. We shall find that the above conventional picture is maintained. The evaporation leads to the same remnant geometry as the process initiated with the gravitational collapse although now the energy of the thunderpop turns out to be independent of the mass of the initial state.

We shall also study a similar process but within a different theory of 1+1 dilaton gravity. It was shown in Ref [8] that the unique theories admitting special conformal symmetries are the models with a classical exponential potential

$$S = \frac{1}{2\pi} \int d^2x \sqrt{-g} \left(R\phi + 4\lambda^2 e^{2\beta\phi} \right) .$$

The classical CGHS action can be recovered for $\beta = 0$ through the field redefinitions $g_{\mu\nu} = \bar{g}_{\mu\nu} e^{-2\bar{\phi}}$, $\phi = e^{-2\bar{\phi}}$. In contrast with the CGHS theory, the black holes arising from the above theory ($\beta \neq 0$) have a non-constant temperature proportional to the mass. Due to the presence of a special symmetry ($g_{\mu\nu} \longrightarrow e^{-\beta\epsilon} g_{\mu\nu}$, $\phi \longrightarrow \phi + \epsilon$), it is possible to construct a semiclassical theory invariant under the above symmetry. So, in this theory one can also find analytic solutions describing the evaporation of a black hole in thermal equilibrium due to the decrease of the incoming thermal flux.

2 Evaporation in the BPP-RST models

The one-parameter family of models [7] interpolating between the BPP and RST models is given by the action

$$S_0 + S_P + \frac{N}{24\pi} \int d^2x \sqrt{-g} \left[(1 - 2a) (\nabla\phi)^2 + (a - 1) \phi R \right] , \quad (2.1)$$

where S_0 is the classical CGHS action, S_P is the Polyakov term and a is an arbitrary real parameter. For $a = \frac{1}{2}$ we recover the RST model and the BPP model

requires $a = 0$. The equations of motion, in conformal gauge, are equivalent to

$$\partial_+ \partial_- (\rho - \phi) = 0, \quad (2.2)$$

$$\partial_+ \partial_- \left(e^{-2\phi} + \frac{Na}{12} \rho \right) + \lambda^2 e^{2(\rho-\phi)} = 0 \quad (2.3)$$

$$\begin{aligned} & e^{-2\phi} \left(4\partial_{\pm} \rho \partial_{\pm} \phi - 2\partial_{\pm}^2 \phi \right) - \frac{N}{12} \left((\partial_{\pm} \rho)^2 - \partial_{\pm}^2 \rho \right) \\ & + \frac{N}{12} \left((2a-1) (\partial_{\pm} \phi)^2 + (a-1) \partial_{\pm}^2 \phi - 2(a-1) \partial_{\pm} \phi \partial_{\pm} \rho \right) \\ & + T_{\pm\pm}^f - \frac{N}{12} t_{\pm} = 0, \end{aligned} \quad (2.4)$$

where $t_{\pm}(x^{\pm})$ are the boundary terms arising from the Polyakov action. In Kruskal gauge ($\rho = \phi$), they reduce to

$$\partial_+ \partial_- \left(e^{-2\phi} + \frac{Na}{12} \phi \right) + \lambda^2 = 0, \quad (2.5)$$

$$\partial_{\pm}^2 \left(e^{-2\phi} + \frac{Na}{12} \phi \right) + T_{\pm\pm}^f - \frac{N}{12} t_{\pm} = 0. \quad (2.6)$$

In the absence of matter ($T_{\pm\pm}^f = 0$), the general solution of (2.5) (2.6) is easily found to be

$$\begin{aligned} e^{-2\phi} + \frac{Na}{12} \phi &= -\lambda^2 x^+ x^- + \frac{N}{12} \int^{x^+} \int^{x^+} t_+(x^+) \\ &+ \frac{N}{12} \int^{x^-} \int^{x^-} t_-(x^-) + bx^+ + cx^- + d, \end{aligned} \quad (2.7)$$

with b, c, d arbitrary constants. Let us consider a solution with $b = c = t_{\pm}(x^{\pm}) = 0$, $d = \frac{M}{\lambda}$

$$e^{-2\phi} + \frac{Na}{12} \phi = -\lambda^2 x^+ x^- + \frac{M}{\lambda}, \quad (2.8)$$

corresponding to a black hole with a curvature singularity (we require $\frac{M}{\lambda} - \alpha > 0$)

$$\alpha = \frac{Na}{24} \left(1 - \log \frac{Na}{24} \right) = -\lambda^2 x^+ x^- + \frac{M}{\lambda}, \quad (2.9)$$

and an apparent horizon ($\partial_+ \phi = 0$) located at

$$x^- = 0. \quad (2.10)$$

If we evaluate the flux of radiation

$$\langle T_{\pm\pm}^f \rangle = -\frac{N}{12} \left((\partial_{\pm} \rho)^2 - \partial_{\pm}^2 \rho + t_{\pm} \right), \quad (2.11)$$

in past and future null infinity using asymptotically flat coordinates ($\lambda x^\pm = \pm e^{\pm\lambda\sigma^\pm}$) we find -due to the anomalous transformation properties of t_\pm - the constant thermal value

$$\langle T_{\sigma^\pm\sigma^\pm}^f \rangle = \frac{N\lambda^2}{48}. \quad (2.12)$$

So that (2.8) represents a radiating black hole in thermal equilibrium with incoming radiation at temperature $T = \frac{\lambda}{2\pi}$ (i.e, the analogue of the Hartle-Hawking state).

Now we shall try to find out a solution without incoming flux which can be continuously matched to (2.8) through the null line $x^+ = x_0^+$. We shall assume that the asymptotically flat coordinate σ^+ is the same as for the initial solution. With this hypothesis we must take $t_{\sigma^+} = 0$ as the no incoming radiation condition, which implies $t_{x^+} = \frac{1}{4(x^+)^2}$ for $x^+ > x_0^+$. Therefore, assuming now the boundary conditions

$$t_{x^+} = \frac{1}{4(x^+)^2} \theta(x^+ - x_0^+), \quad (2.13)$$

$$t_{x^-} = 0, \quad (2.14)$$

as the new input for (2.7) and demanding continuity to the metric and its derivative through $x^+ = x_0^+$, we obtain (for $x^+ > x_0^+$) the dynamical solution ($k = \frac{N}{12}$)

$$e^{-2\phi} + \frac{Na}{12}\phi = -\lambda^2 x^+ (x^- + \Delta) - \frac{k}{4} \left(\log \frac{x^+}{x_0^+} + 1 \right) + \frac{M}{\lambda}, \quad (2.15)$$

where $\Delta = -\frac{k}{4\lambda^2 x_0^+}$.

The new asymptotically flat coordinates $\{\tilde{\sigma}^+, \tilde{\sigma}^-\}$ are

$$\lambda x^+ = e^{\lambda\tilde{\sigma}^+}, \quad (2.16)$$

$$-\lambda(x^- + \Delta) = e^{-\lambda\tilde{\sigma}^-}. \quad (2.17)$$

The equality of σ^+ and $\tilde{\sigma}^+$ confirms our previous assumption and thus guarantees that solution (2.15) has vanishing incoming radiation.

The curvature singularity of (2.15) is

$$\alpha = \frac{ka}{2} \left(1 - \log \frac{ka}{2} \right) = -\lambda^2 x^+ (x^- + \Delta) - \frac{k}{4} \left(\log \frac{x^+}{x_0^+} + 1 \right) + \frac{M}{\lambda}, \quad (2.18)$$

whereas the apparent horizon

$$-\lambda^2 x^+ (x^- + \Delta) = \frac{k}{4}, \quad (2.19)$$

recedes as the black hole evaporates and intersects the curvature singularity at the point (see Fig.1)

$$x_{int}^+ = x_0^+ e^{\frac{4}{k}(\frac{M}{\lambda} - \alpha)}, \quad (2.20)$$

$$x_{int}^- = \frac{k}{4\lambda^2 x_0^+} \left[1 - e^{-\frac{4}{k}(\frac{M}{\lambda} - \alpha)} \right]. \quad (2.21)$$

We notice that the evaporation time differs from that obtained in the process originated by gravitational collapse [7].

Following the conventional reasoning [5] [6] we shall try to implement the cosmic censorship conjecture by matching the evaporating black hole solution (2.15) with a stable (not-radiating) solution on the null line $x^- = x_{int}^-$. From [7] we know that the stable solutions of (2.5) (2.6) are

$$e^{-2\phi} + \frac{Na}{12}\phi = -\lambda^2 x^+ (x^- + \Delta) - \frac{k}{4} \log \left(-\lambda^2 x^+ (x^- + \Delta) \right) + C, \quad (2.22)$$

corresponding to the boundary conditions $t_{x^\pm} = \frac{1}{4(x^\pm)^2}$. Imposing continuity to the metric through $x^- = x_{int}^-$ we obtain the value of the constant C

$$C = \frac{k}{4} \left(\log \frac{k}{4} - 1 \right) + \alpha. \quad (2.23)$$

The remnant geometry (2.22) coincides exactly with the final state resulting from the evaporation of a black hole formed by collapse of a shock wave, as discussed in [7]. However, in the present process the energy of the thunderpop emanating from the intersecting point is independent of the initial state and takes the limiting value $E_{thunderpop} = -\frac{\lambda k}{4}$.

3 Evaporation in the exponential models

In this section we shall investigate the evaporation process of black holes within a different theory but with boundary conditions similar to that considered in the previous section. Our starting point is the semiclassical action of the exponential models [8]

$$S = \frac{1}{2\pi} \int d^2x \sqrt{-g} \left(\phi R + 4\lambda^2 e^{\beta\phi} - \frac{1}{2} \sum_{i=1}^N (\nabla f_i)^2 \right) + S_P + \frac{N\beta}{96\pi} \int d^2x \sqrt{-g} \left(\phi R + \beta (\nabla\phi)^2 \right), \quad (3.1)$$

where S_P is the Polyakov action. The local counterterm of (3.1) can be obtained from the Polyakov effective action respect to the metric invariant under the special conformal symmetry. This way one ensures the invariance of the semiclassical action and so the existence of a free field. The equations of motion of (3.1), in conformal gauge, are

$$\partial_+ \partial_- \left(\rho - \frac{\beta}{2} \phi \right) = 0, \quad (3.2)$$

$$\partial_+ \partial_- \left[\left(1 + \frac{N\beta}{24} \right) \phi + \frac{N\rho}{12} \right] + \lambda^2 e^{(2\rho + \beta\phi)} = 0, \quad (3.3)$$

$$\begin{aligned}
& - \left(1 + \frac{N\beta}{12}\right) \partial_{\pm}^2 \phi + 2 \left(1 + \frac{N\beta}{12}\right) \partial_{\pm} \phi \partial_{\pm} \rho \\
& + \frac{N}{12} \left[(\partial_{\pm} \rho)^2 - \partial_{\pm}^2 \rho + \frac{\beta^2}{4} (\partial_{\pm} \phi)^2 + t_{\pm} (x^{\pm}) \right] = T_{\pm\pm}^f. \quad (3.4)
\end{aligned}$$

In Kruskal gauge ($2\rho = \beta\phi$) these equations can be rewritten as

$$\partial_+ \partial_- 2\rho = -\lambda^2 \beta \gamma e^{4\rho}, \quad (3.5)$$

$$e^{2\rho} \partial_{\pm}^2 e^{-2\rho} = \beta \gamma \left(T_{\pm\pm}^f - \frac{N}{12} t_{\pm} \right) \quad (3.6)$$

where $\gamma = \frac{1}{1 + \frac{N\beta}{12}}$.

We see that 2ρ satisfies the Liouville equation (3.5). Its general solution can be expressed as

$$\rho = \frac{1}{4} \log \frac{\partial_+ A \partial_- B}{(1 + \lambda^2 \beta \gamma AB)^2}, \quad (3.7)$$

with $A(B)$ an arbitrary function of the x^+ (x^-) coordinate. Inserting (3.7) into the constrained equations (3.6) we obtain Riccati differential equations with respect to $\log \partial_+ A$ and $\log \partial_- B$.

$$\frac{1}{4} (\partial_+ \log \partial_+ A)^2 - \frac{1}{2} \partial_+^2 \log \partial_+ A = \beta \gamma \left(T_{++}^f - \frac{N}{12} t_+ \right), \quad (3.8)$$

$$\frac{1}{4} (\partial_- \log \partial_- B)^2 - \frac{1}{2} \partial_-^2 \log \partial_- B = \beta \gamma \left(T_{--}^f - \frac{N}{12} t_- \right). \quad (3.9)$$

Taking $T_{\pm\pm}^f = t_{\pm} = 0$, one obtains the solution

$$ds^2 = \frac{-dx^+ dx^-}{\frac{\lambda^2 \beta \gamma}{C} + C x^+ x^-}. \quad (3.10)$$

For $C < 0$ (and $\lambda^2 \beta < 0$) [8] the above metric represents a black hole with a singularity curve

$$\frac{\lambda^2 \beta \gamma}{C} + C x^+ x^- = 0, \quad (3.11)$$

and an apparent horizon located at

$$x^- = 0. \quad (3.12)$$

Proceeding in a similar way as in the past section we can calculate the flux of ingoing and outgoing radiation in the asymptotically flat coordinates $\{\sigma^{\pm}\}$

$$\pm \sqrt{|C|} x^{\pm} = e^{\pm \sqrt{|C|} \sigma^{\pm}}, \quad (3.13)$$

getting

$$\langle T_{\sigma^{\pm} \sigma^{\pm}}^f \rangle = \frac{N|C|}{48}, \quad (3.14)$$

which corresponds to a radiating black hole in thermal equilibrium with ingoing radiation at temperature $T = \frac{\sqrt{|C|}}{2\pi}$. The black hole mass can be calculated

immediately [8]: $M = \frac{2}{\beta\pi}\sqrt{|C|}$. Therefore, the Hawking temperature turns out to be proportional to the mass: $T = \frac{\beta}{4}M$. The existence of 2D black holes whose mass is proportional to their temperature was noted in Ref. [9]. At this point, the next step would be to solve equations (3.8) and (3.9) with the modified boundary conditions (2.13) (2.14) and impose continuity to the resulting metric and its derivative through $x^+ = x_0^+$. However this leads to an excessively complicated analytical treatment. In order to simplify the analysis we shall admit that, in addition to the discontinuity in the boundary function t_{x^+} , there is an incoming shock-wave at $x^+ = x_0^+$

$$T_{++}^f = \frac{1}{2x_0^+\beta\gamma}\delta(x^+ - x_0^+). \quad (3.15)$$

With this assumption, and in the limit $N\beta \gg 1$, the solutions of (3.8) and (3.9), for $x^+ > x_0^+$, lead to the metric

$$ds^2 = -\frac{\left(\frac{x_0^+}{x^+}\right)^{\frac{1}{2}}}{\frac{\lambda^2\beta\gamma}{C} + Cx_0^+x^- \left(\log\frac{x^+}{x_0^+} + 1\right)} dx^+ dx^-. \quad (3.16)$$

The most relevant feature of this solution is that the singularity curve

$$\frac{\lambda^2\beta\gamma}{C} + Cx_0^+x^- \left(\log\frac{x^+}{x_0^+} + 1\right) = 0, \quad (3.17)$$

and the horizon

$$\frac{\lambda^2\beta\gamma}{C} + Cx_0^+x^- \left(\log\frac{x^+}{x_0^+} + 1\right) + 2Cx_0^+x^- = 0, \quad (3.18)$$

do not intersect each other at any finite point (see Fig.2). Moreover, in contrast with the BPP-RST models the relation of the asymptotically flat coordinate $\tilde{\sigma}^+$ of (3.16) and the Kruskal one x^+ , which regulates the incoming flux, is not the same as for the initial solution (3.10). In fact, in past null infinity, this relation becomes

$$\frac{dx^+}{\sqrt{|C|} \left(x_0^+x^+\right)^{\frac{1}{2}} \left(\log\frac{x^+}{x_0^+} + 1\right)} = d\tilde{\sigma}^+. \quad (3.19)$$

This tells us that the thermal bath has not been actually removed. Using (3.19) we can calculate the value of $\langle T_{\tilde{\sigma}^+\tilde{\sigma}^+}^f \rangle$ in past null infinity for $x^+ > x_0^+$

$$\begin{aligned} \langle T_{\tilde{\sigma}^+\tilde{\sigma}^+}^f \rangle &= -\frac{N}{12} \left[\frac{|C|x_0^+}{4(x^+)} \left(\log\frac{x^+}{x_0^+} + 1\right)^2 - \frac{1}{8(x^+)^2} \right. \\ &\quad \left. - \frac{3}{16(x^+)^2} \left(\frac{\log\frac{x^+}{x_0^+} + 3}{\log\frac{x^+}{x_0^+} + 1}\right)^2 \right]. \end{aligned} \quad (3.20)$$

We see that the flux goes to zero at late times ($x^+ \rightarrow \infty$) whereas at $x^+ = x_0^+$

$$\langle T_{\bar{\sigma}^+\bar{\sigma}^+}^f \rangle |_{x_0^+} = -\frac{N}{12} \left[\frac{|C|}{4} - \frac{29}{16(x_0^+)^2} \right]. \quad (3.21)$$

At this point it seems natural to choose the point x_0^+ which makes (3.21) to coincide with the constant value of $\langle T_{\sigma^+\sigma^+}^f \rangle = \frac{N}{48}|C|$ of the initial solution (i.e, $x_0^+ = \sqrt{\frac{29}{8|C|}}$). In this way the external flux starts to decrease at $x^+ = x_0^+$ (see Fig.3), vanishes at retarded time x_1^+ and for $x^+ > x_1^+$ becomes a small negative energy flux vanishing again at infinity.

In conclusion, the semiclassical solution implies that the black hole evaporation never ceases. One can argue that the non-vanishing incoming flux for $x^+ > x_0^+$ prevents the complete evaporation[§]. However after $x^+ > x_1^+$ the flux becomes negative and produces the opposite effect and, despite of this, the apparent horizon never meets the singularity. This qualitative behaviour of the solution can be understood in terms of thermodynamic arguments. In the evaporation process the temperature slows down because it is proportional to the mass and the third law of thermodynamics [10] does not permit to achieve the absolute zero temperature (i.e, the complete evaporation).

After completing this work we were informed that the black hole solution (3.10) also appears in Ref. [11].

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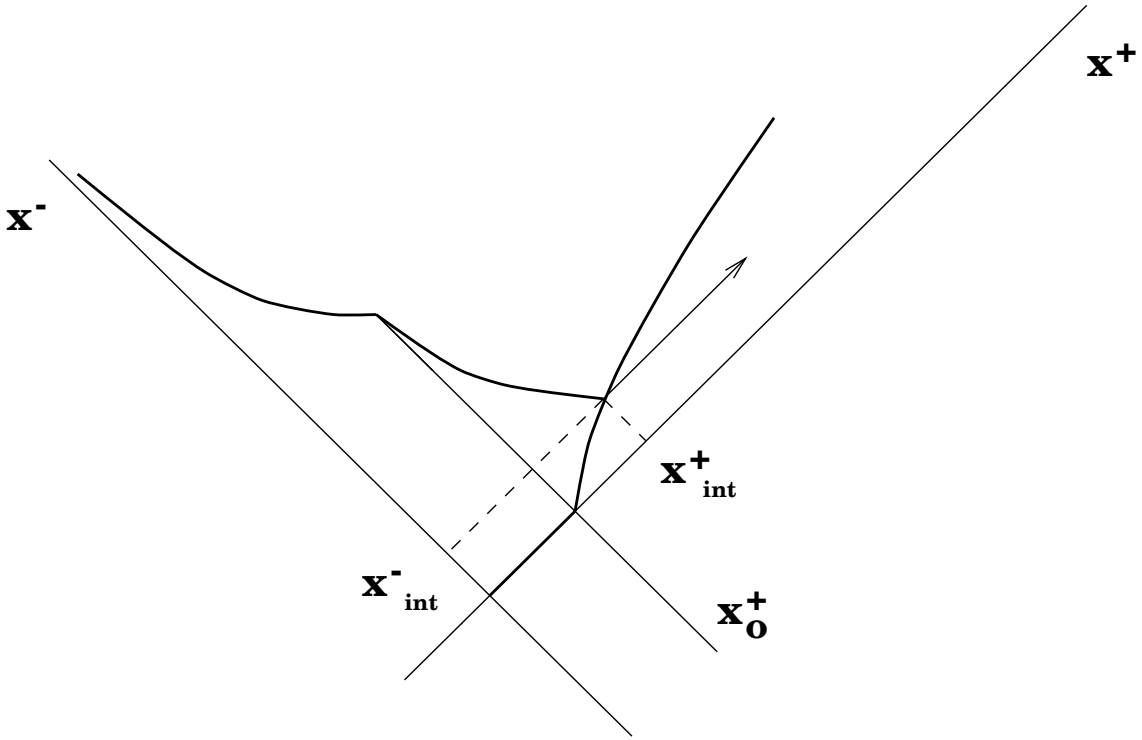


Fig.1 Kruskal diagram of an evaporating black hole in the BPP-RST models

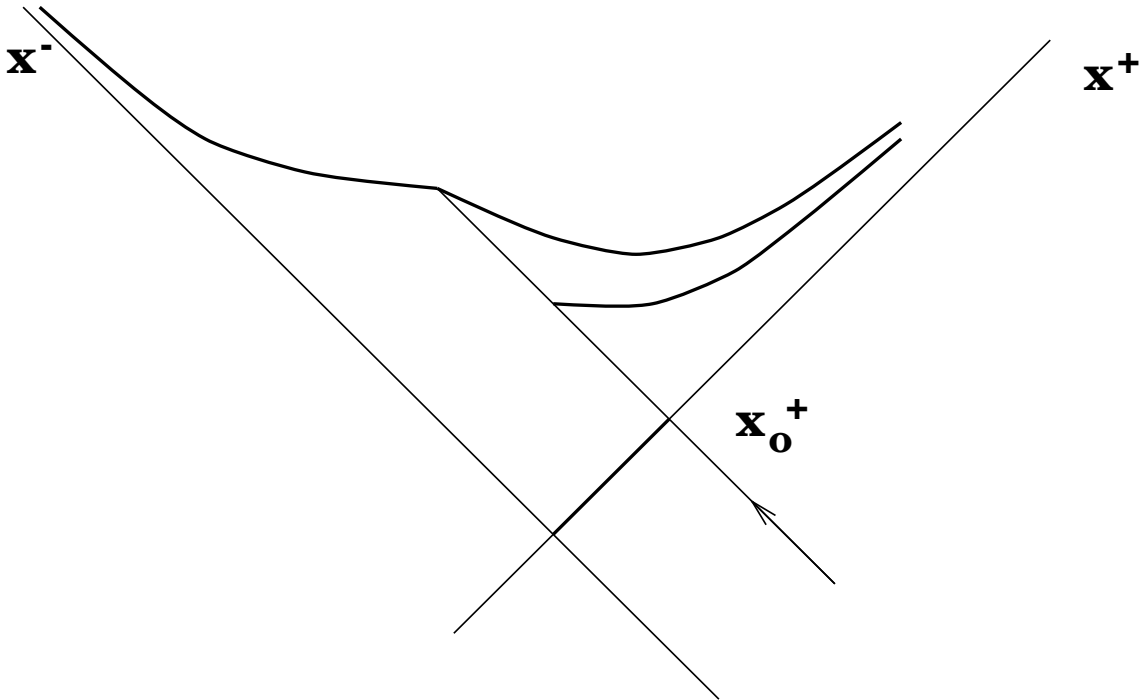


Fig.2 Kruskal diagram of an evaporating black hole in the exponential model

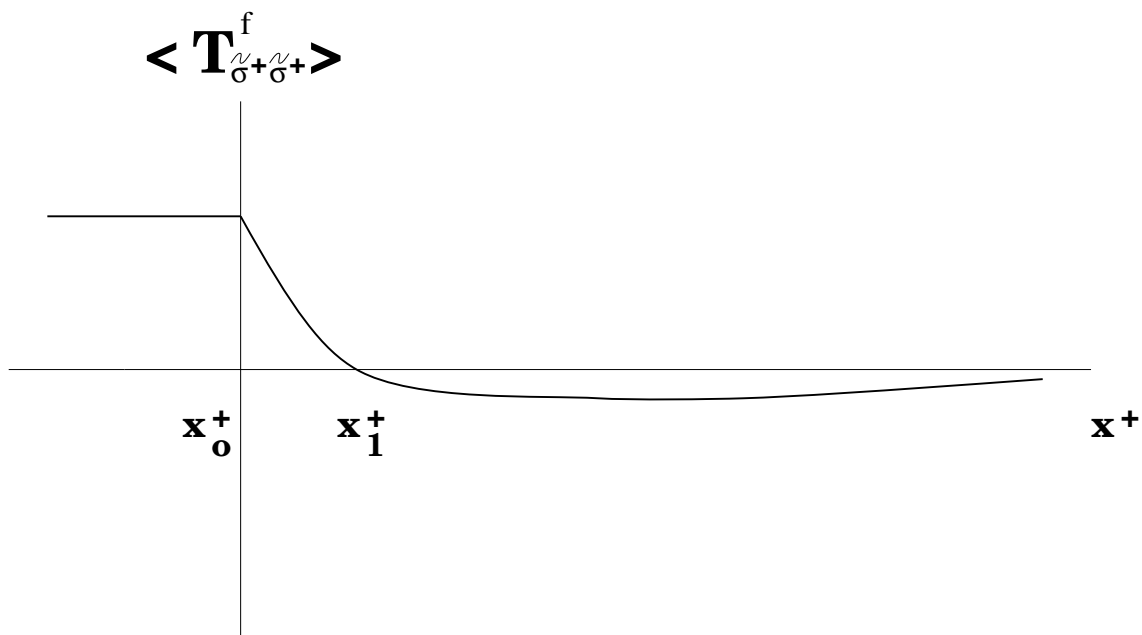


Fig.3 Flux of incoming radiation associated with the boundary condition $t_{x^+} = \frac{1}{4(x^+)^2} \theta(x^+ - x_0^+)$ in the exponential model